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Neutron star mergers and the high-density equation of state

Xiamen-CUSTIPEN Workshop on the EOS of dense neutron-rich matter in the era of gravitational wave astronomy

Xiamen, 06/01/2019

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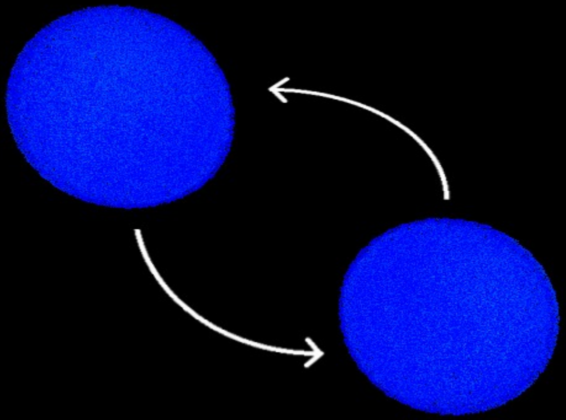
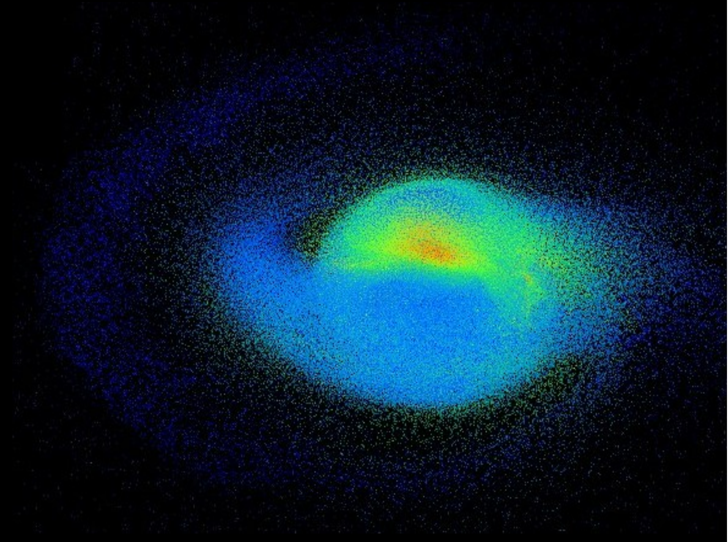
(GSI Darmstadt)

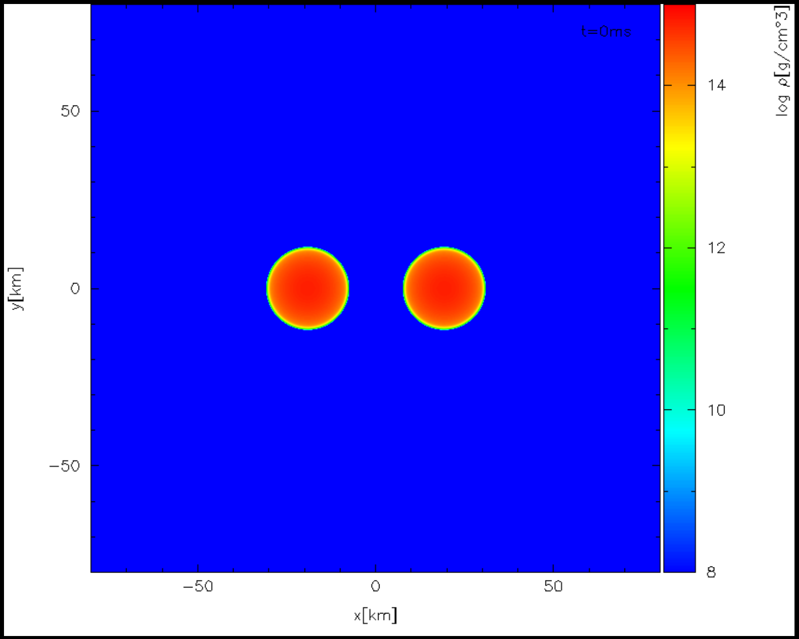




Outline

- Multi-messenger interpretation of GW170817 → lower limit on NS radii

  → Collapse behavior (EoS dependence of BH formation)
- Signatures of the QCD phase transition

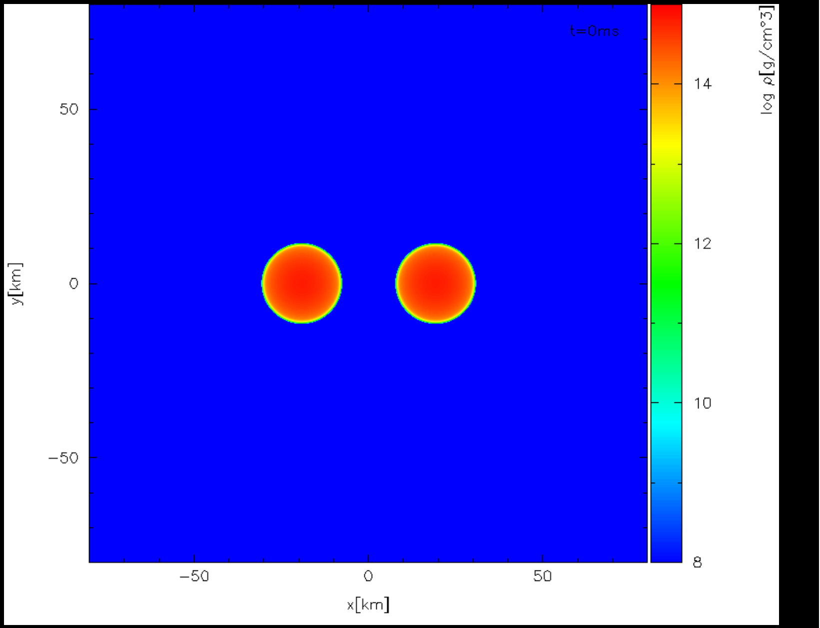


$$M_{\rm tot} = 3.4\ M_\odot$$

←

$$M_{\rm tot} = 3.5\ M_\odot$$

→



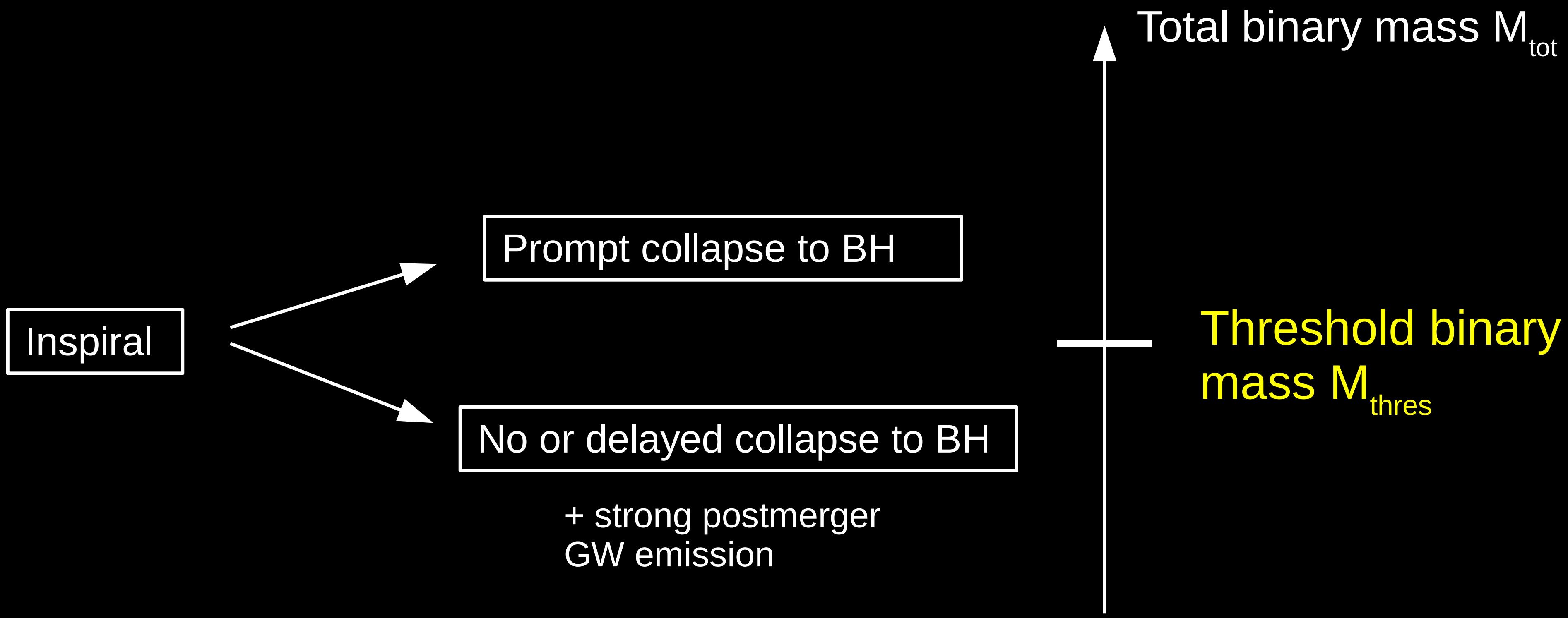
Shen EoS

$$\longrightarrow \quad M_{\rm thres} = (3.45 \pm 0.05)\ \ M_\odot$$ (for this particular EoS)

Collapse behavior: Prompt vs. delayed (/no) BH formation

Relevant for: EoS constraints through M_{max} measurement, Conditions for short GRBs, Mass ejection, Electromagnetic counterparts powered by thermal emission, NS radius constraints !!!

Collapse behavior



Total binary mass M_{tot}

Prompt collapse to BH

Inspiral

Threshold binary mass M_{thres}

No or delayed collapse to BH

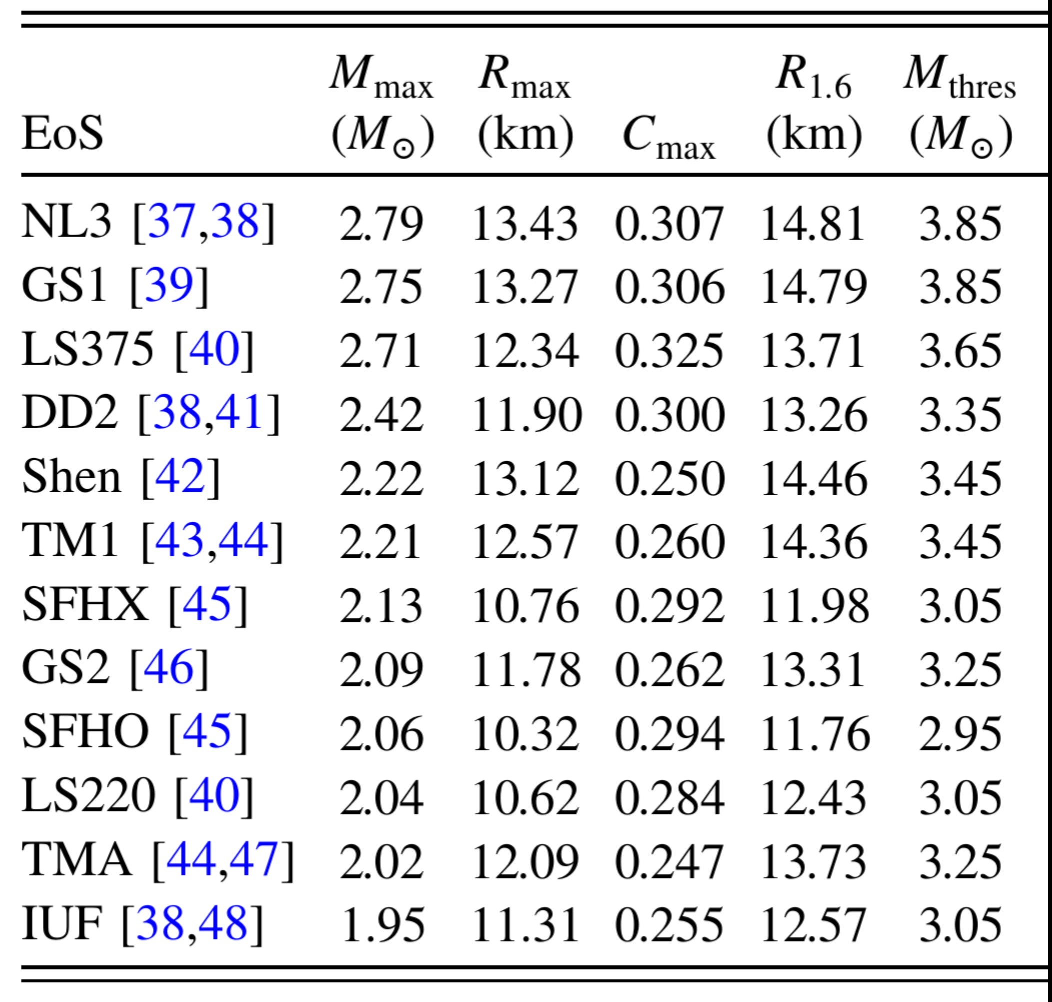
+ strong postmerger GW emission

EoS dependent - somehow M_{max} should play a role

Simulations reveal M_{thres}

TOV properties of nonrotating stars, i.e. EoS characteristics

Merger property from simulations



EoS	$M_{\max}$ ($M_\odot$)	$R_{\max}$ (km)	$C_{\max}$	$R_{1.6}$ (km)	M_{thres} ($M_\odot$)
NL3 [37,38]	2.79	13.43	0.307	14.81	3.85
GS1 [39]	2.75	13.27	0.306	14.79	3.85
LS375 [40]	2.71	12.34	0.325	13.71	3.65
DD2 [38,41]	2.42	11.90	0.300	13.26	3.35
Shen [42]	2.22	13.12	0.250	14.46	3.45
TM1 [43,44]	2.21	12.57	0.260	14.36	3.45
SFHX [45]	2.13	10.76	0.292	11.98	3.05
GS2 [46]	2.09	11.78	0.262	13.31	3.25
SFHO [45]	2.06	10.32	0.294	11.76	2.95
LS220 [40]	2.04	10.62	0.284	12.43	3.05
TMA [44,47]	2.02	12.09	0.247	13.73	3.25
IUF [38,48]	1.95	11.31	0.255	12.57	3.05

Bauswein et al. 2013

Smooth particle hydrodynamics + conformal flatness

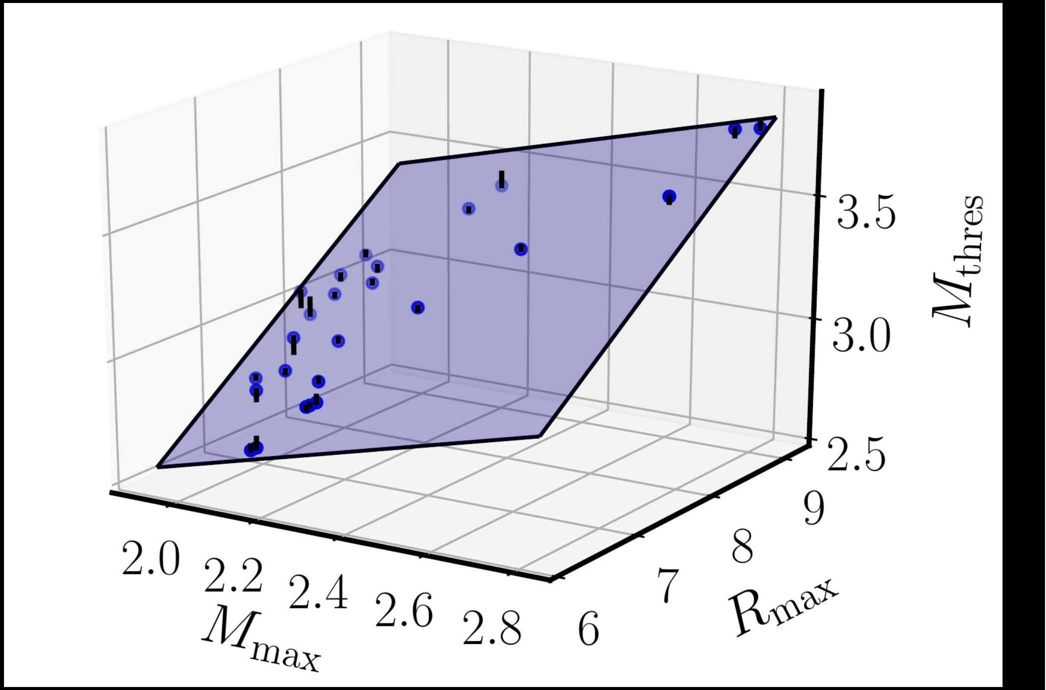
Threshold binary mass

- Empirical relation from simulations with different M_{tot} and EoS
- Fits (to good accuracy):

$$M_{\mathrm{thres}} = M_{\mathrm{thres}}(M_{\mathrm{max}}, R_{\mathrm{max}}) = \left(-3.38 \frac{G M_{\mathrm{max}}}{c^2 R_{\mathrm{max}}} + 2.43\right) M_{\mathrm{max}}$$

$$M_{\mathrm{thres}} = M_{\mathrm{thres}}(M_{\mathrm{max}}, R_{1.6}) = \left(-3.6 \frac{G\, M_{\mathrm{max}}}{c^2\, R_{1.6}} + 2.38\right) M_{\mathrm{max}}$$

- Both better than 0.06 M_{sun}



EoS constraints from GW170817*

→ lower bound on NS radii

(recall: upper bound from tidal deformability)

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* See also Margalit & Metzger 2017, Shibata et al. 2017, Rezzolla et al. 2018, Radice et al. 2018, Ruiz & Shapiro 2018, ... for other EoS constraints in the context of GW170817

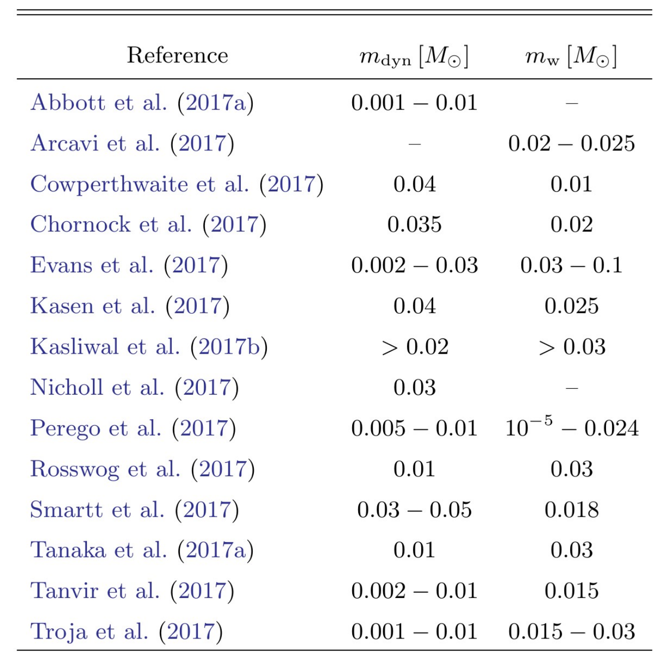
A simple but robust NS radius constraint from GW170817

- High ejecta mass inferred from electromagnetic transient

  (high compared to simulations)

  → provides strong support for a delayed/no collapse in GW170817

  → even asymmetric mergers that directly collapse do not produce such massive ejecta



Reference	$m_{\rm dyn}$ $[M_\odot]$	$m_{\rm w}$ $[M_\odot]$
Abbott et al. (2017a)	0.001 – 0.01	–
Arcavi et al. (2017)	–	0.02 – 0.025
Cowperthwaite et al. (2017)	0.04	0.01
Chornock et al. (2017)	0.035	0.02
Evans et al. (2017)	0.002 – 0.03	0.03 – 0.1
Kasen et al. (2017)	0.04	0.025
Kasliwal et al. (2017b)	> 0.02	> 0.03
Nicholl et al. (2017)	0.03	–
Perego et al. (2017)	0.005 – 0.01	10^{-5} – 0.024
Rosswog et al. (2017)	0.01	0.03
Smartt et al. (2017)	0.03 – 0.05	0.018
Tanaka et al. (2017a)	0.01	0.03
Tanvir et al. (2017)	0.002 – 0.01	0.015
Troja et al. (2017)	0.001 – 0.01	0.015 – 0.03

Compilation in Cote et al 2018

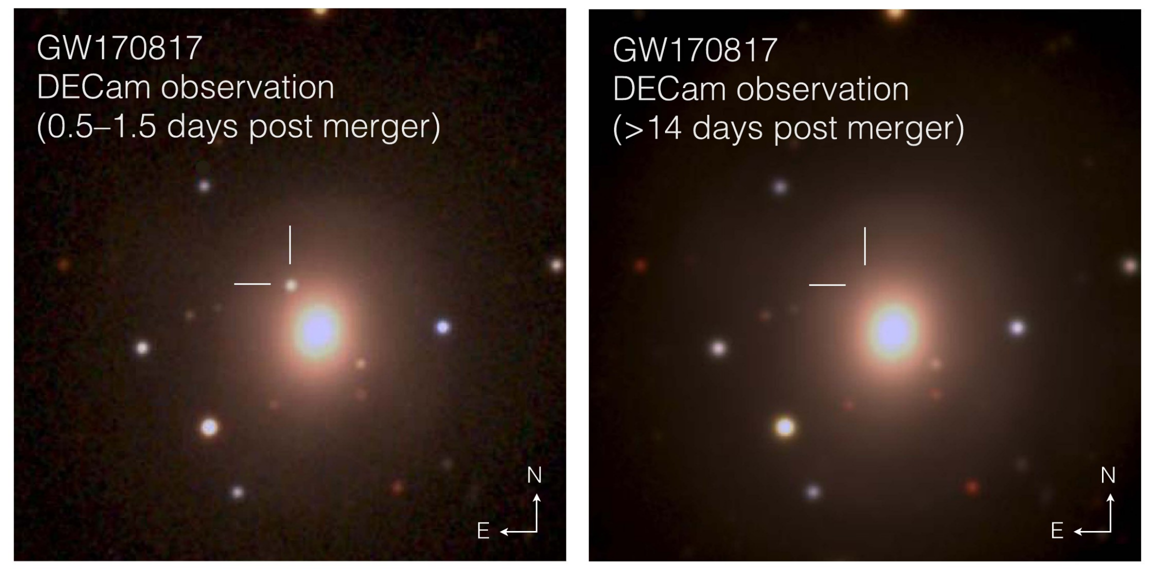
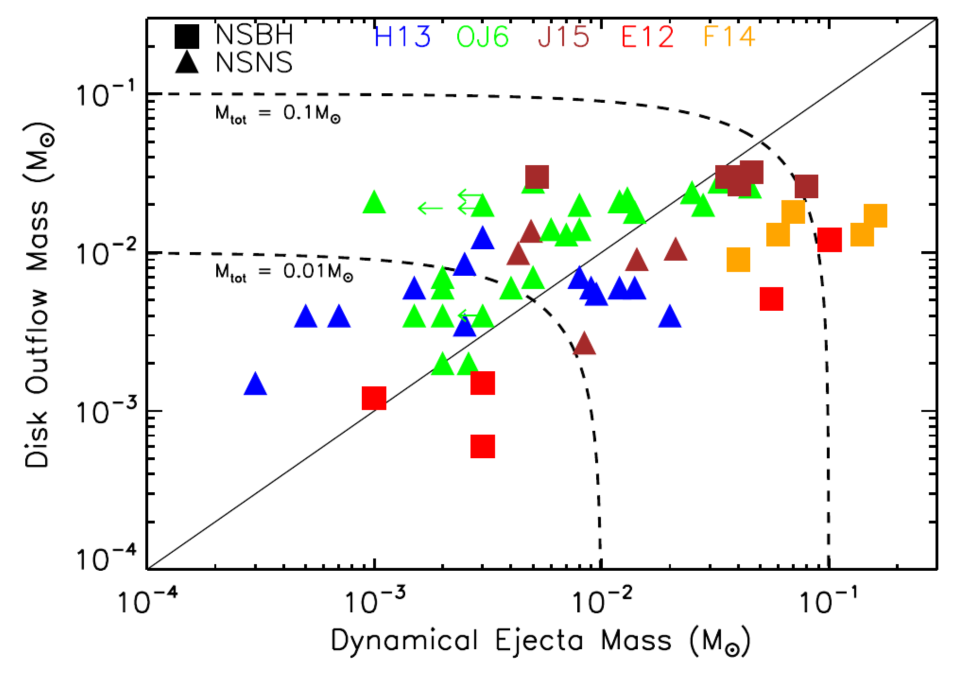


Figure 1. NGC4993 *grz* color composites ($1'.5 \times 1'.5$). Left: composite of detection images, including the discovery *z* image taken on 2017 August 18 00:05:23 UT and the g and r images taken 1 day later; the optical counterpart of GW170817 is at R.A., decl. =197.450374, −23.381495. Right: the same area two weeks later.

Soares-Santos et al 2017

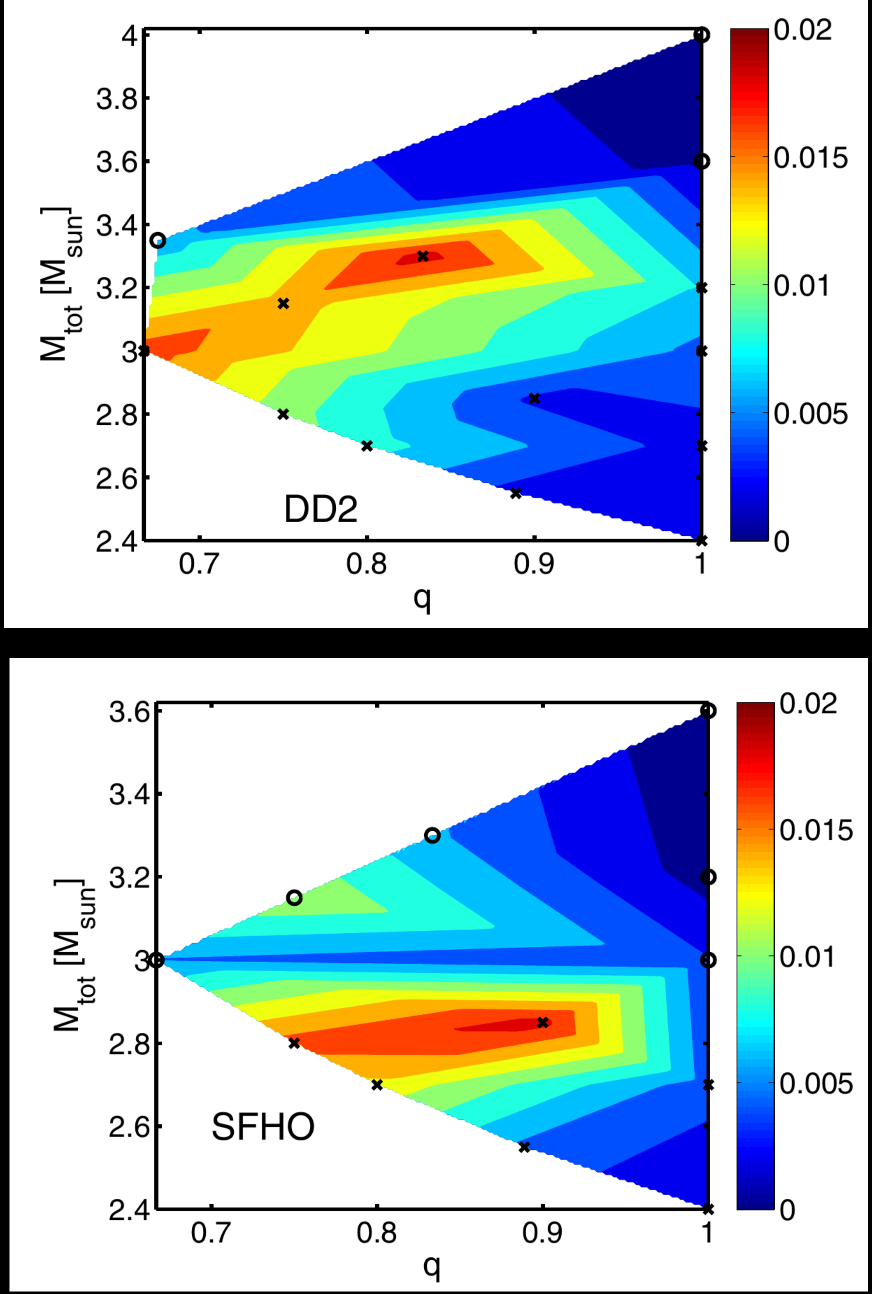
- Ejecta masses depend on EoS and binary masses
- Note: high mass points already to soft EoS (tentatively/qualitatively)
- Prompt collapse leads to reduced ejecta mass
- Light curve depends on ejecta mass:

→ 0.02 - 0.05 M_{sun} point to delayed collapse



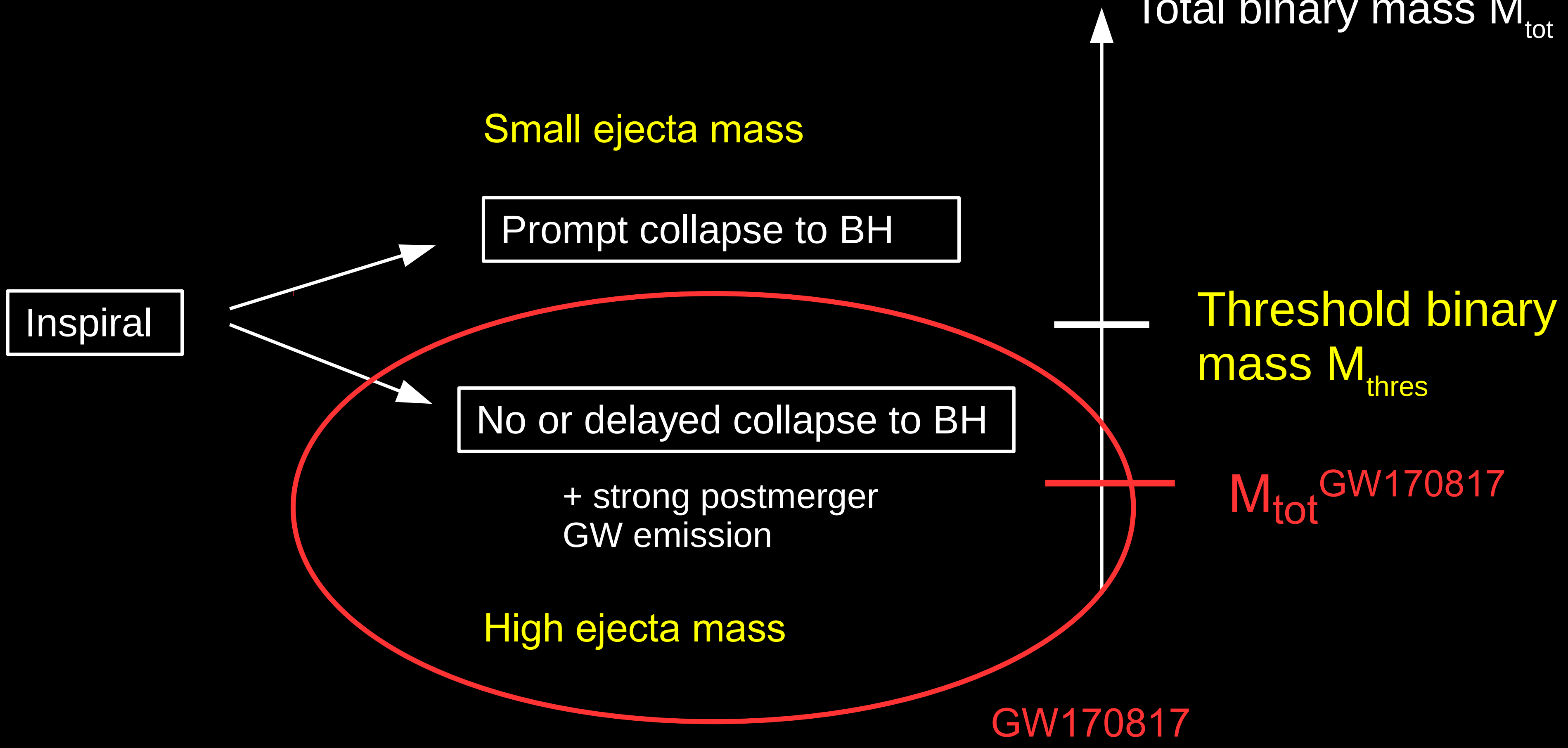
Compilation Wu et al 2016: dynamical and secular ejecta comparable

Only dynamical ejecta



Bauswein et al. 2013

Collapse behavior



Inspiral

Small ejecta mass

Prompt collapse to BH

No or delayed collapse to BH

+ strong postmerger
GW emission

High ejecta mass

GW170817

Total binary mass M_{tot}

Threshold binary mass M_{thres}

$M_{tot}^{GW170817}$

(1) If GW170817 was a delayed (/no) collapse:

$$M_{\mathrm{thres}} > M_{\mathrm{tot}}^{GW170817}$$

(2) Recall: empirical relation for threshold binary mass for prompt collapse:

$$M_{\mathrm{thres}} = \left(-3.38\frac{G\,M_{\mathrm{max}}}{c^2\,R_{\mathrm{max}}} + 2.43\right)\,M_{\mathrm{max}} > 2.74\;M_{\odot}$$

(with M_{max}, R_{max} unknown)

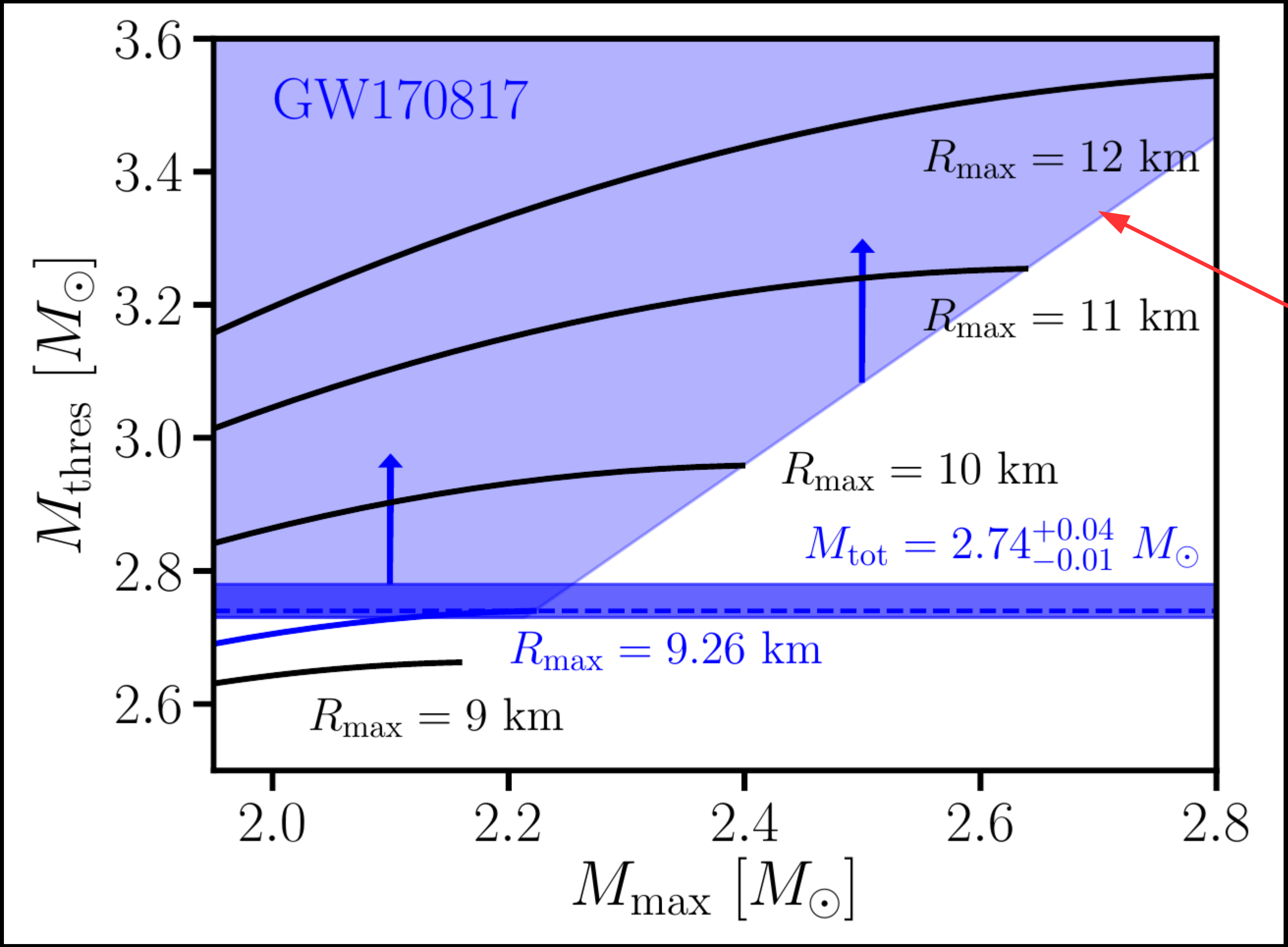
(3) Causality: speed of sound $v_S \leq c$ $\Rightarrow M_{\mathrm{max}} \leq \frac{1}{2.82}\frac{c^2\,R_{\mathrm{max}}}{G}$

- Putting things together:

$$M_{\mathrm{tot}}^{GW170817} \leq \left(-3.38\frac{G\,M_{\mathrm{max}}}{c^2\,R_{\mathrm{max}}} + 2.43\right) M_{\mathrm{max}} \leq \left(-\frac{3.38}{2.82} + 2.43\right)\frac{1}{2.82}\frac{c^2\,R_{\mathrm{max}}}{G}$$

→ Lower limit on NS radius

Bauswein et al. 2017

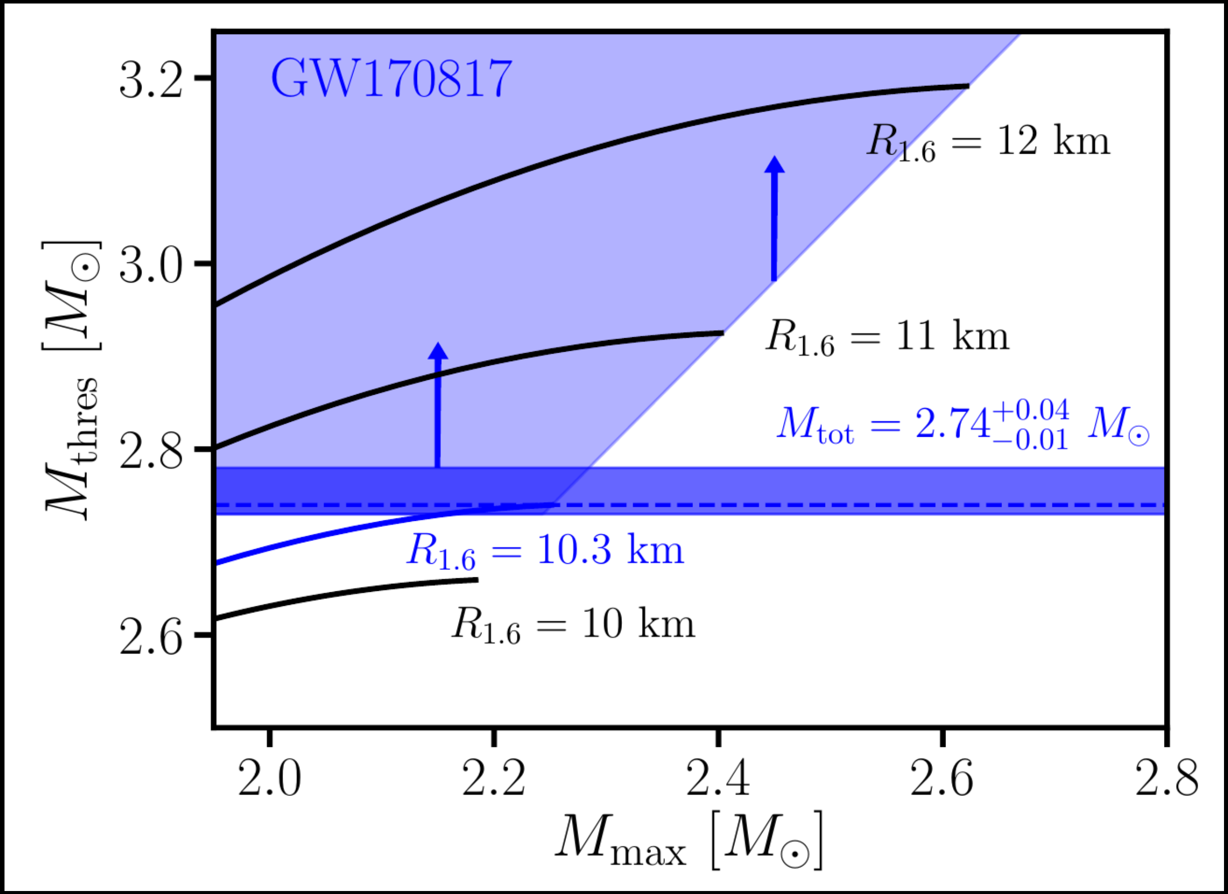


Constraint by causality

$$M_{\rm thres} = \left(-3.38\frac{GM_{\max}}{c^2R_{\max}} + 2.43\right)M_{\max}$$

$$M_{\rm thres} \geq 1.2M_{\max}$$

Bauswein et al. 2017

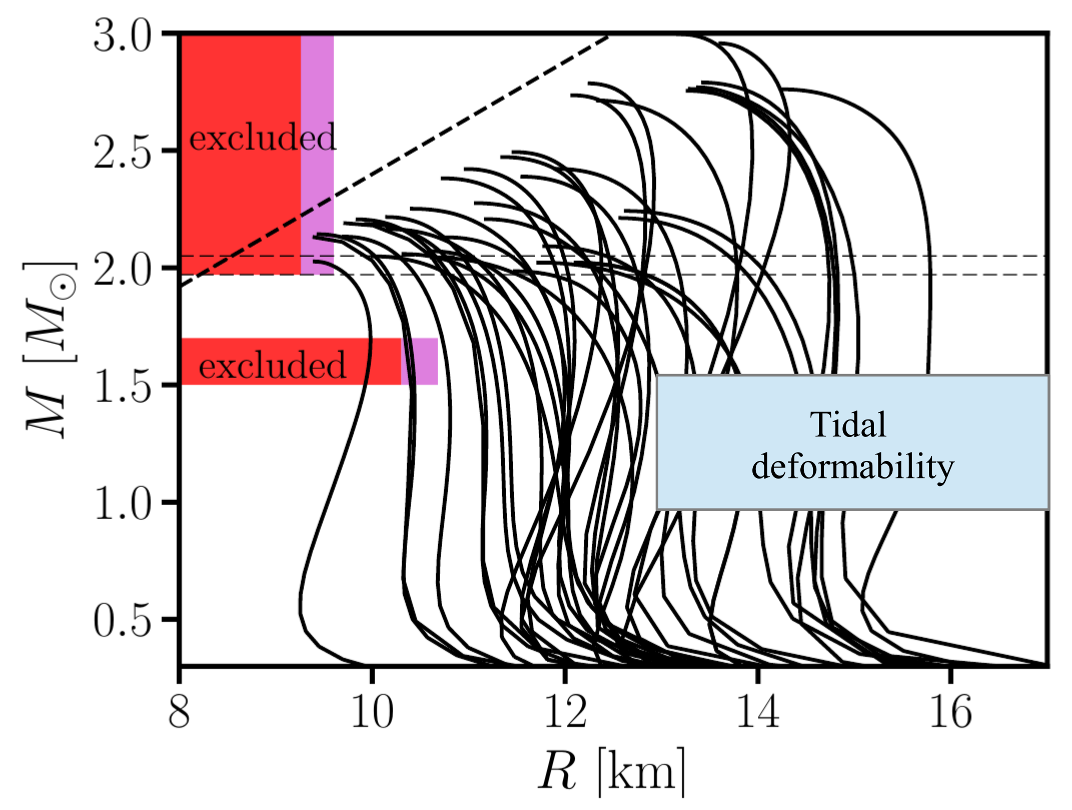


$$M_{\rm thres} = \left(-3.6\frac{G\,M_{\rm max}}{c^2\,R_{1.6}} + 2.38\right)\,M_{\rm max}$$

$$v_S = \sqrt{\frac{dP}{de}} \leq c \;\; \rightarrow \;\; M_{\rm max} \leq \kappa R_{1.6} \;\; \Rightarrow \;\; M_{\rm thres} \geq 1.2 M_{\rm max}$$

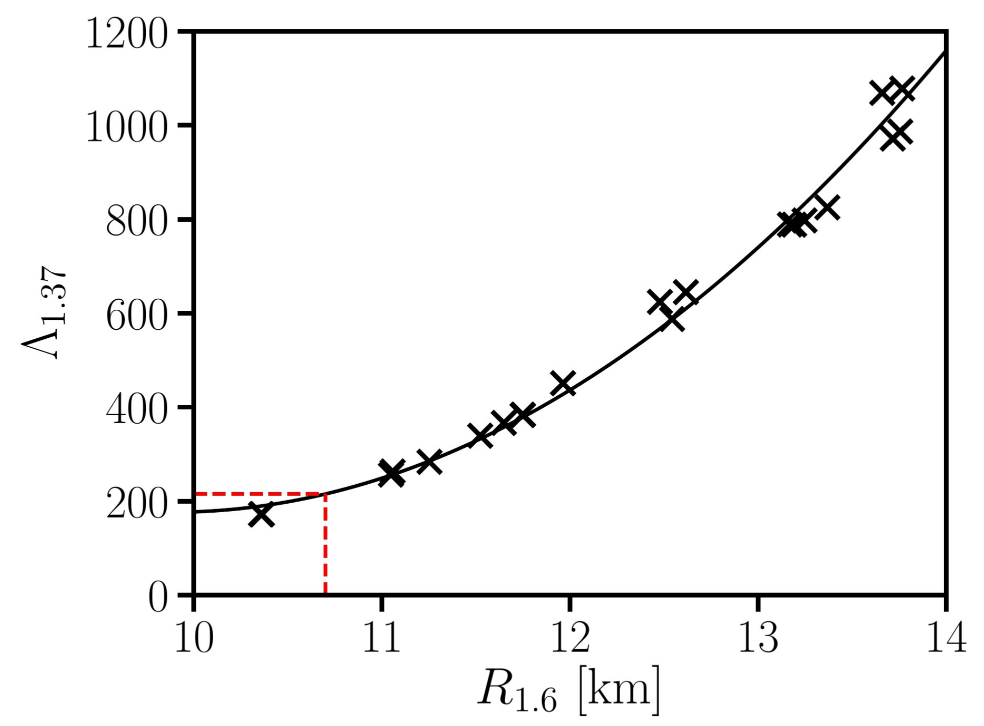
NS radius constraint from GW170817

- $R_{max} > 9.6$ km
- $R_{1.6} > 10.7$ km
- Excludes very soft nuclear matter



Bauswein et al. 2017

Radius vs. tidal deformability



Bauswein, unpubl.

- Radius and tidal deformability scale tightly → **Lambda > 210**
- Note: limit cannot be much larger because there are EoS (with somewhat larger Lambda) that do NOT result in a prompt collapse !!

  (argument turned around – full EoS coverage essential !!!)

  (little bit of buffer because we were very conservative in Bauswein et al 2017)

Discussion - robustness

- Binary masses well measured with high confidence error bar
- Clearly defined working hypothesis: delayed collapse

  → testable by refined emission models

  → as more events are observed more robust distinction
- Very conservative estimate, errors can be quantified
- Empirical relation can be tested by more elaborated simulations (but unlikely that MHD or neutrinos can have strong impact on M_{thres})
- Confirmed by semi-analytic collapse model
- Low-SNR constraint !!!

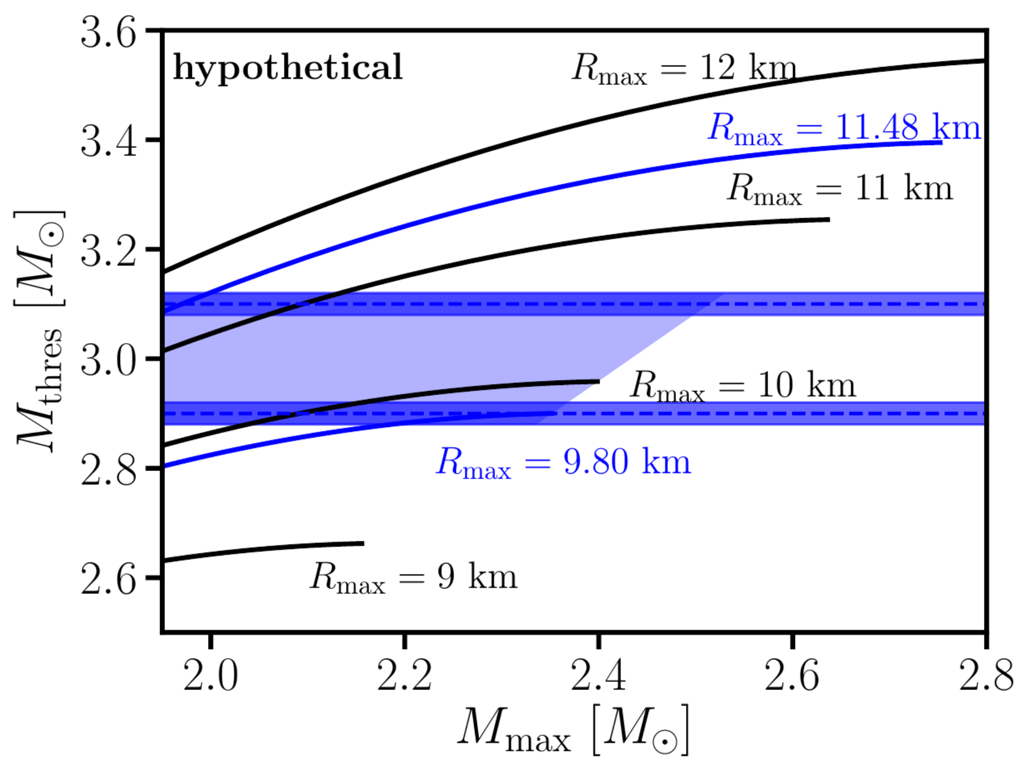
Future

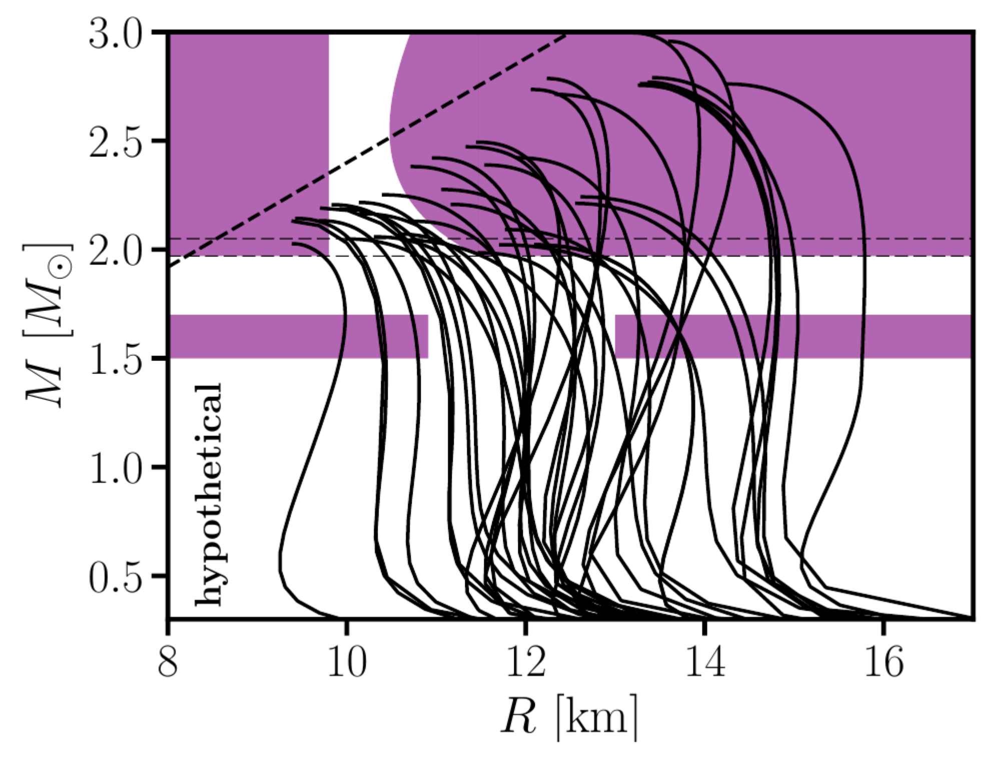
- Any new detection can be employed if it allows distinction between prompt/delayed collapse
- With more events in the future our comprehension of em counterparts will grow → more robust discrimination of prompt/delayed collapse events
- Low-SNR detections sufficient !!! → that's the potential for the future

  → we don't need louder events, but more

  → complimentary to existing ideas for EoS constraints

Future detections (hypothetical discussion)





→ as more events are observed, bands converge to true $M_{\rm thres}$

→ prompt collapse constrains $M_{\max}$ from above

Bauswein et al. 2017

Future: Maximum mass

- Empirical relation

$$M_{\mathrm{thres}} = \left(-3.6\frac{G\,M_{\mathrm{max}}}{c^2\,R_{1.6}} + 2.38\right)\,M_{\mathrm{max}}$$

- Sooner or later we'll know $R_{1.6}$ (e.g. from postmerger) and M_{thres} (from several events – through presense/absence of postmerger GW emission or em counterpart)

=> direct inversion to get precise estimate of M_{max}

(see also current estimates e.g. by Margalit & Metzger, Rezzolla et al, Ruiz & Shapiro, Shibata et al., ...)

Postmerger GW emission*

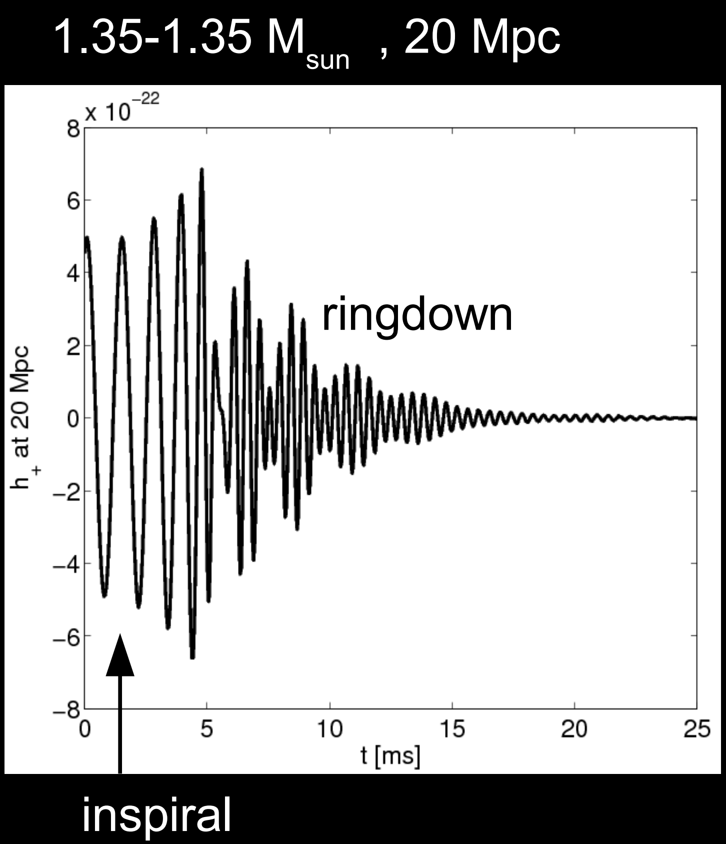
(dominant frequency of postmerger phase)

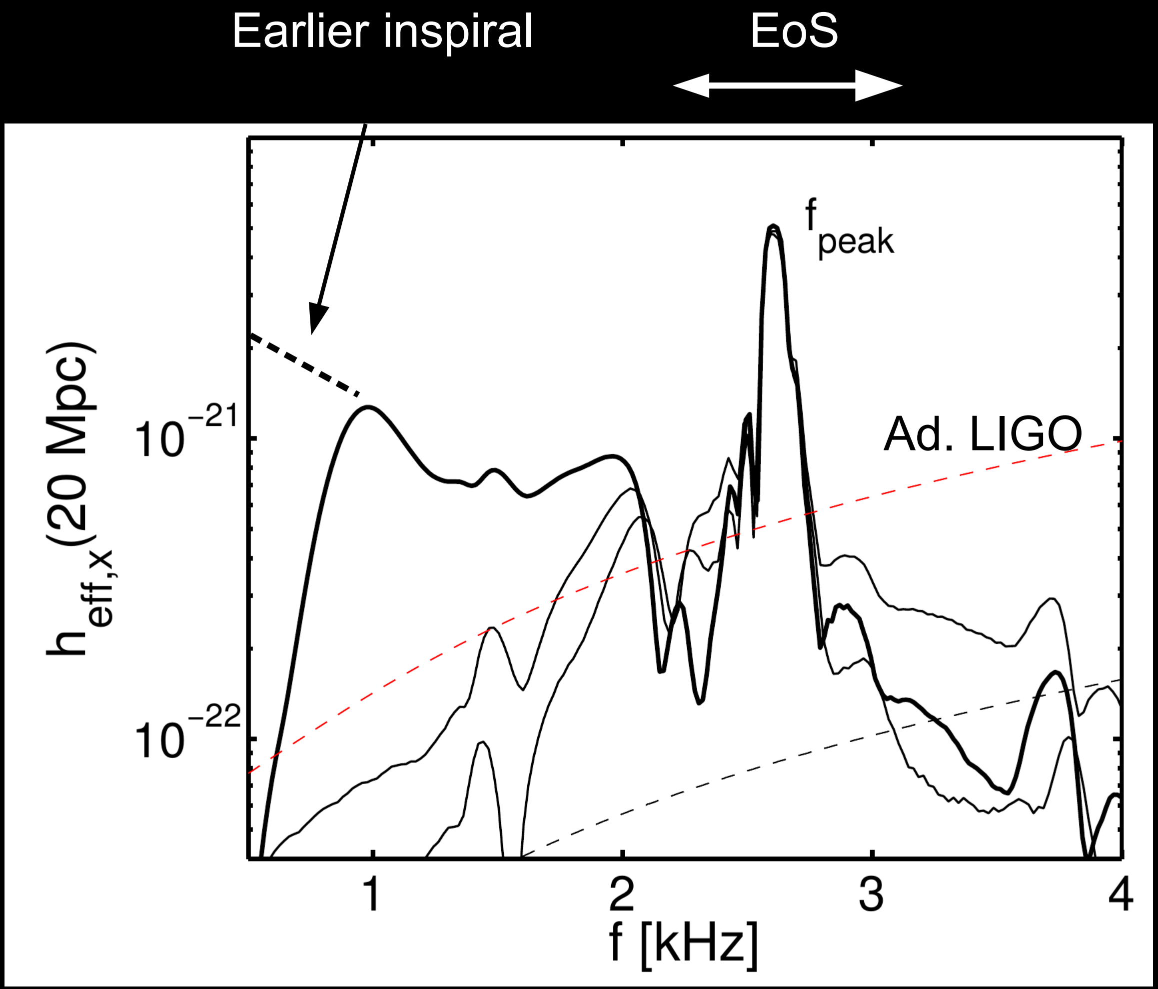
→ determine properties of EoS/NSs

→ complementary to inspiral

* not detected for GW170817 – expected for current sensitivity and d=40 Mpc (Abbott et al. 2017)

Postmerger

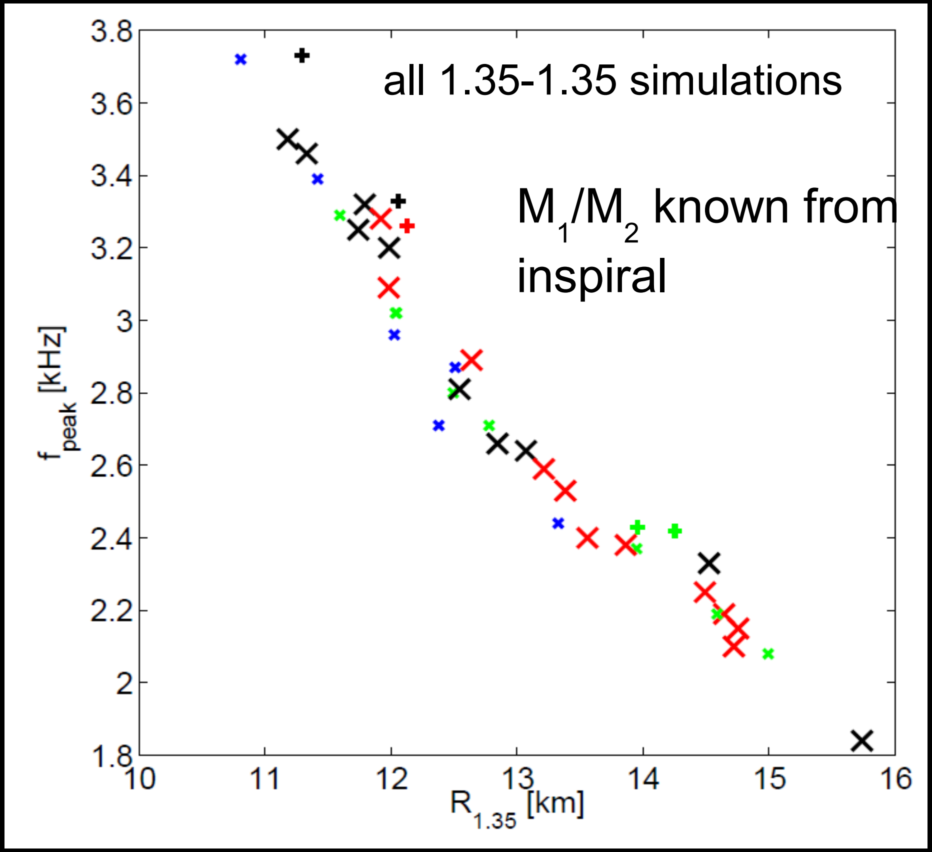

1.35-1.35 M_{sun} , 20 Mpc

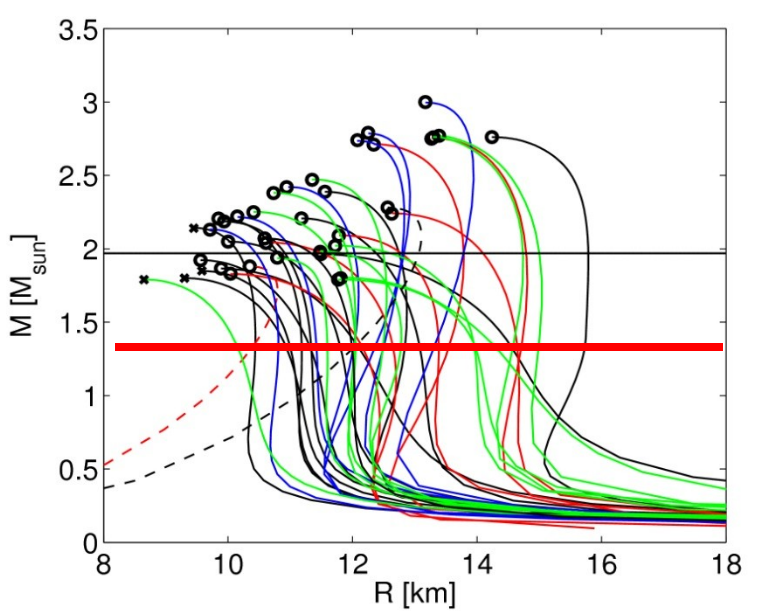


Dominant postmerger oscillation frequency f_{peak}

Very characteristic (robust feature in all models)

Gravitational waves – EoS survey





characterize EoS by radius of nonrotating NS with 1.35 M_{sun}

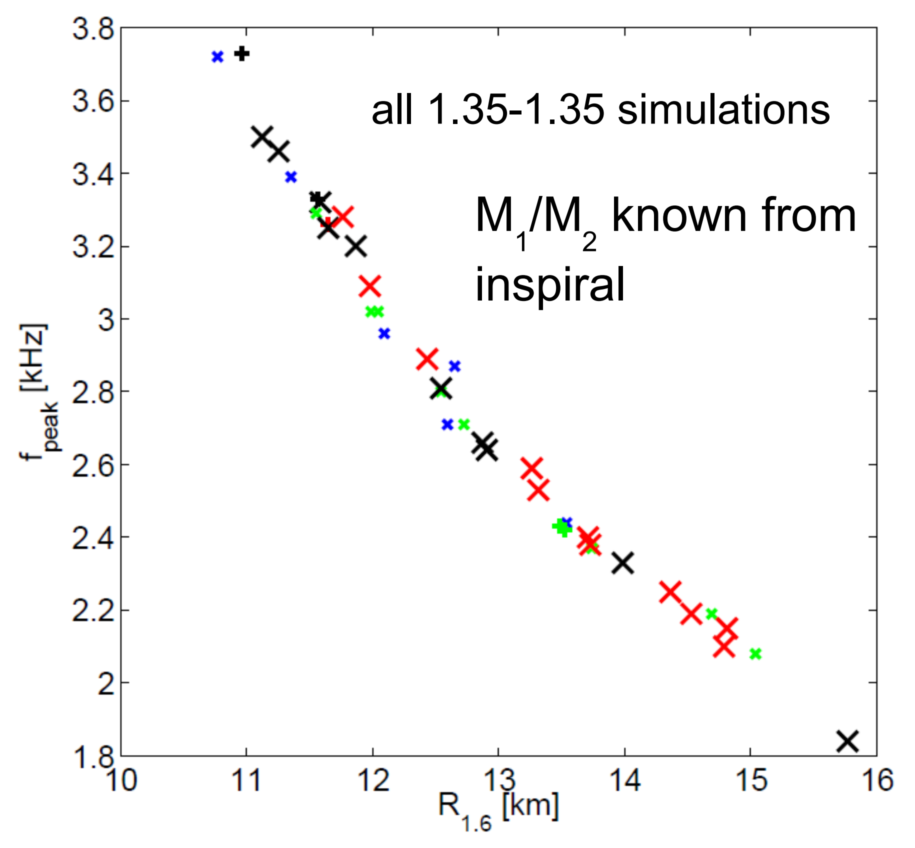
Bauswein et al. 2012

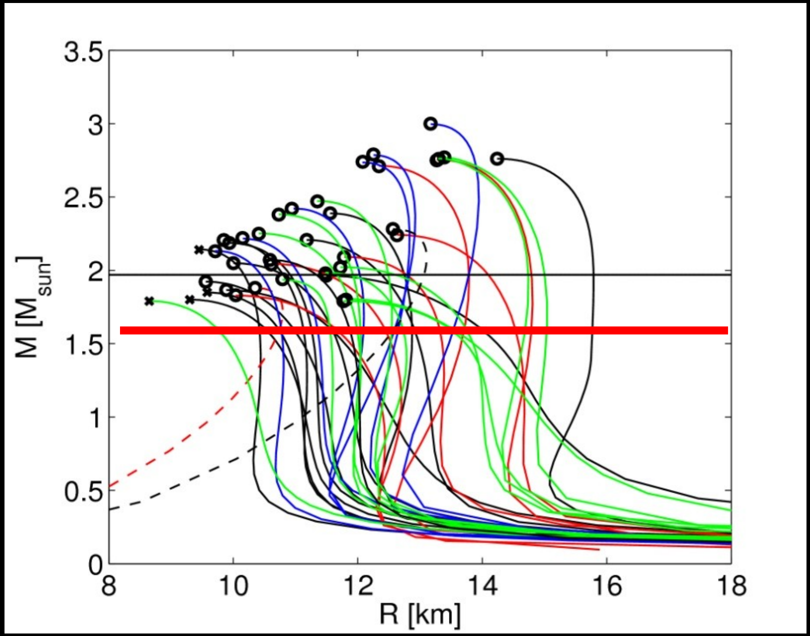
Pure TOV/EoS property => Radius measurement via f_{peak}

Here only 1.35-1.35 Msun mergers (binary masses measurable) – similar relations exist for other fixed binary setups !!!

~ 40 different NS EoSs

Gravitational waves – EoS survey





characterize EoS by radius of nonrotating NS with 1.6 M_{sun}

Bauswein et al. 2012

Pure TOV/EoS property => Radius measurement via f_{peak}

Smaller scatter in empirical relation (< 200 m) → smaller error in radius measurement

Note: R of 1.6 M_{sun} NS scales with f_{peak} from 1.35-1.35 M_{sun} mergers (density regimes comparable)

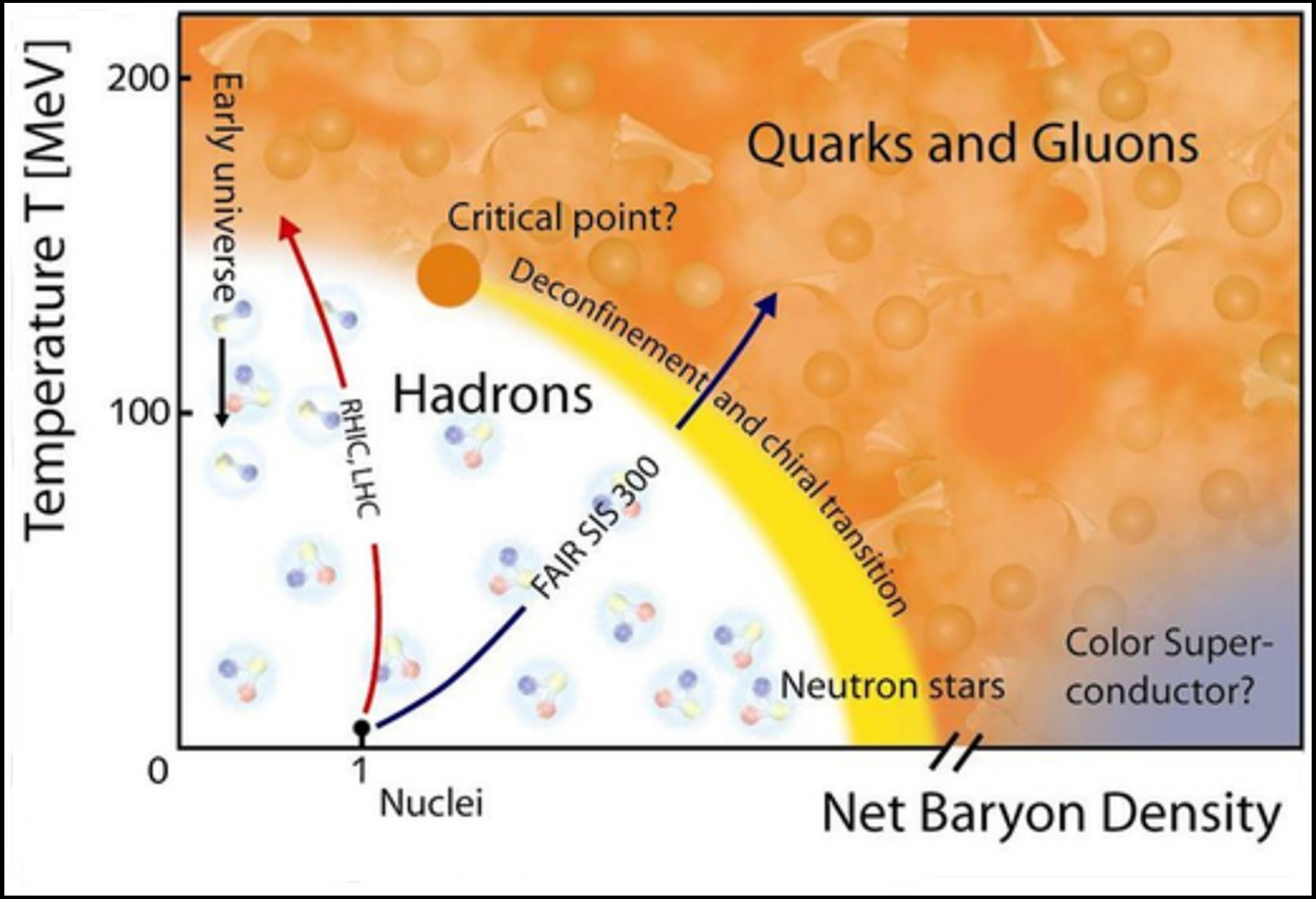
GW data analysis: Clark et al 2014, Clark et al 2016, Chatziioannou et al 2017, …

→ detectable at a few 10 Mpc

Observable signature of (QCD) phase transition

Phase diagram of matter

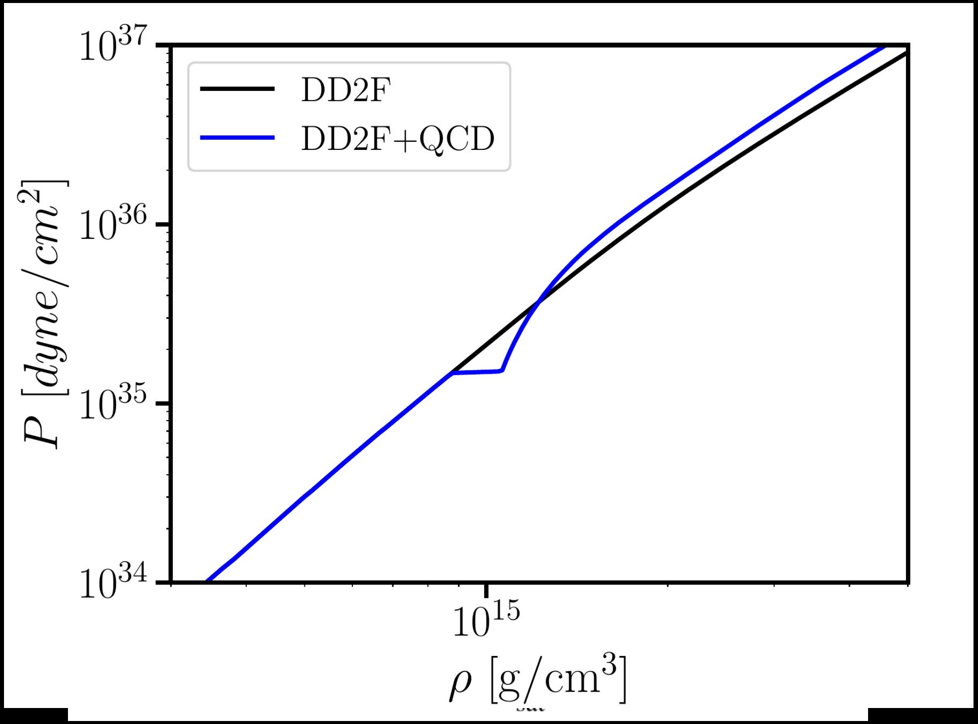
GSI/FAIR

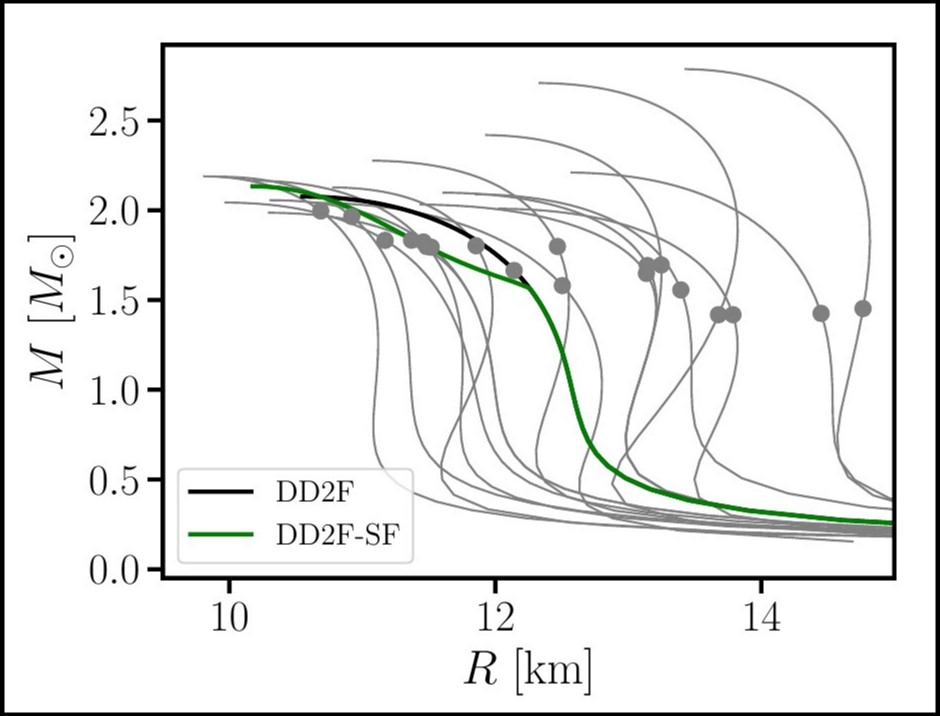


Does the phase transition to quark-gluon plasma occur (already) in neutron stars or only at higher densities ?

EoS with 1st-order phase transition to quark matter

Bauswein et al. 2018

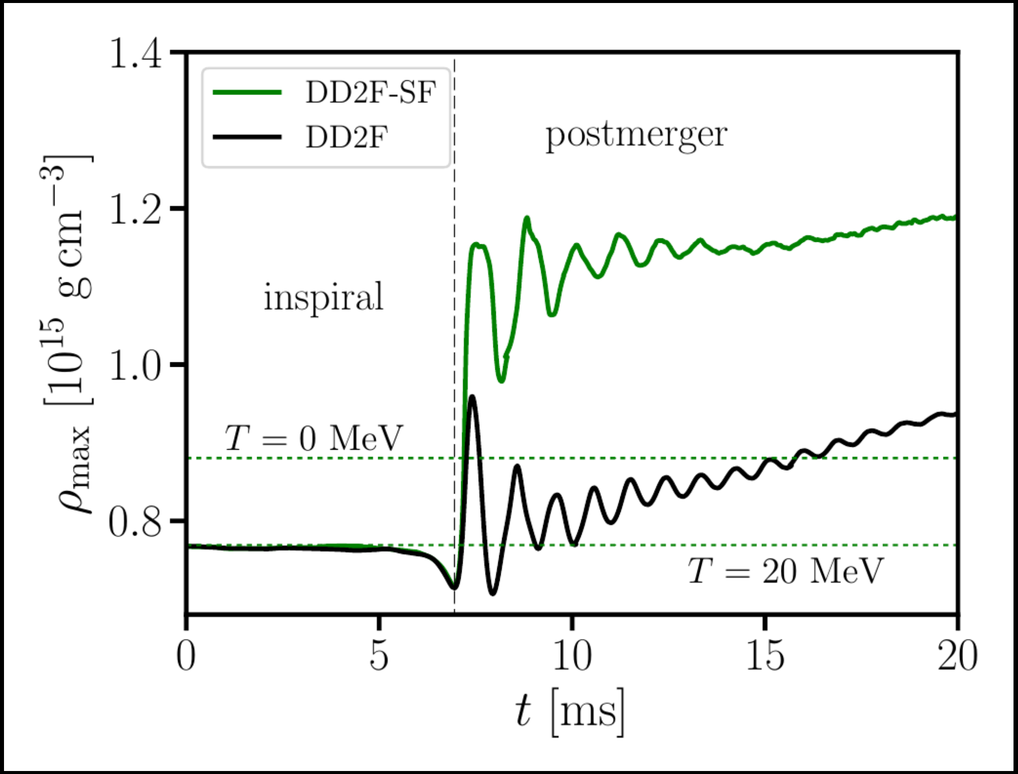


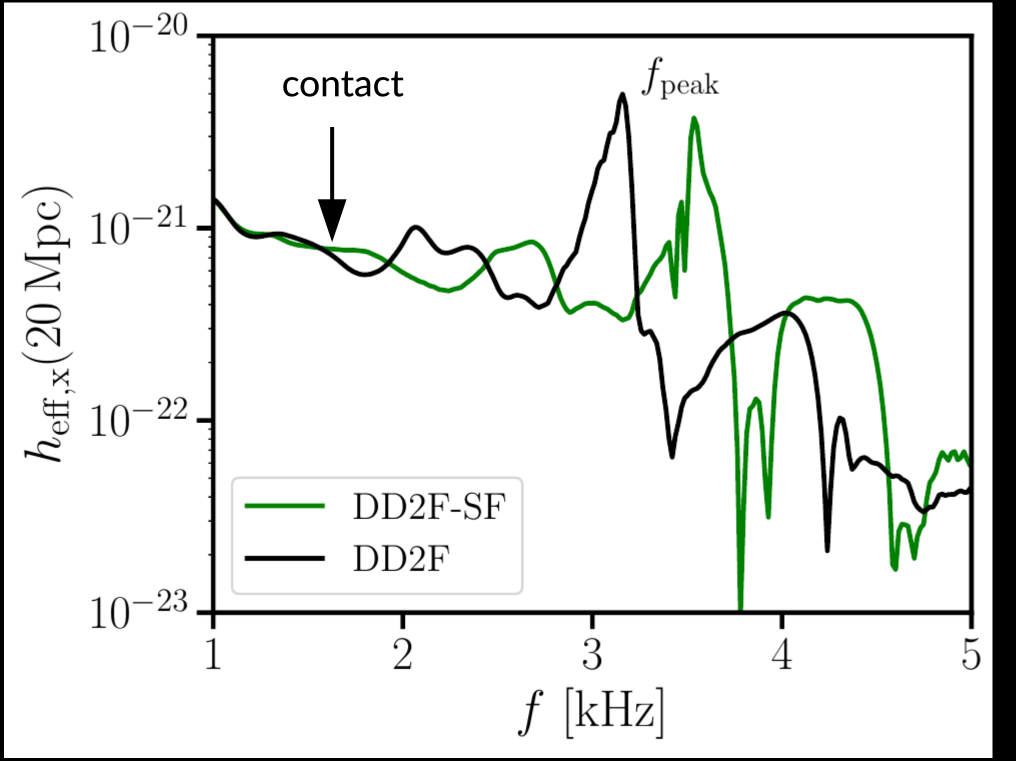


- EoS from Fischer et al. 2018 – as one example for an EoS with a strong 1st-order phase transition to deconfined quarks
- Difficult to measure transition in mergers through inspiral: Lambda very small, high mass star probably less frequent

Merger simulations

- GW spectrum 1.35-1.35 Msun





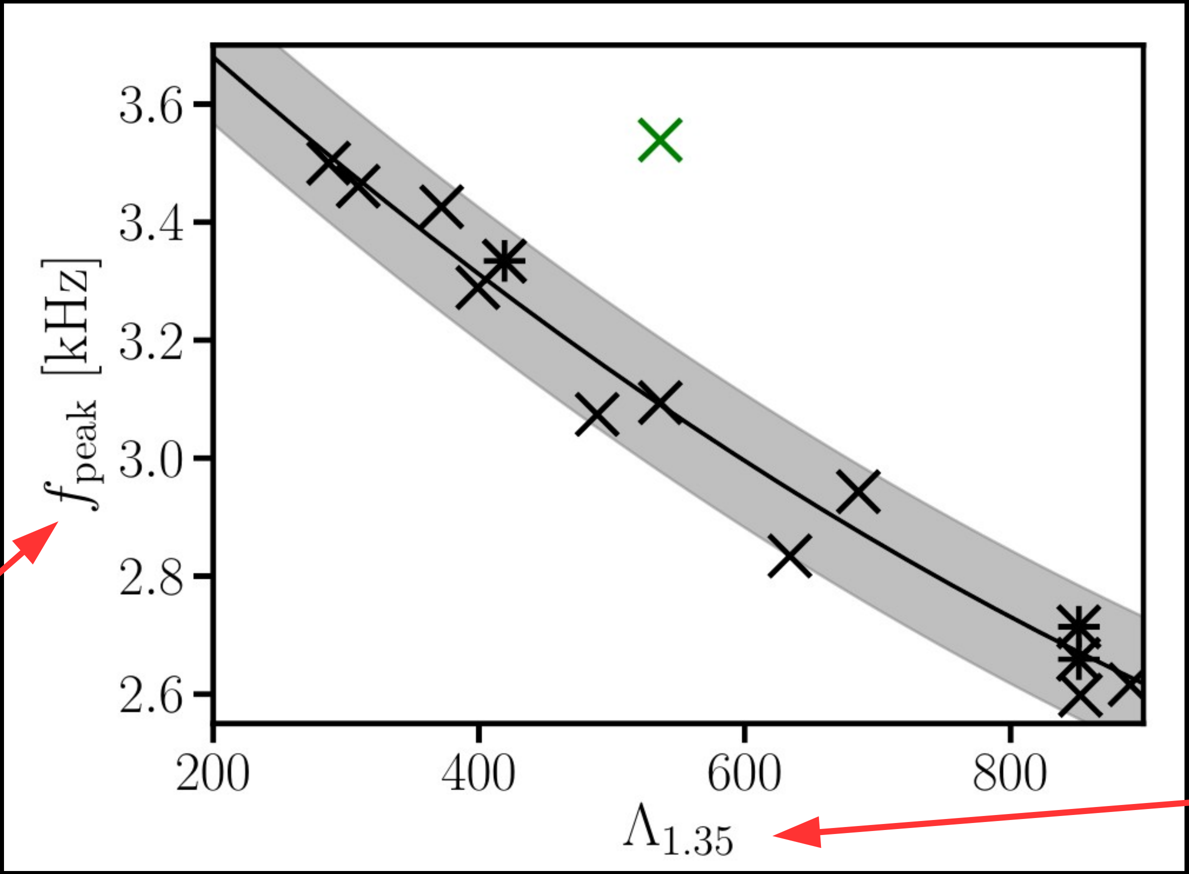
Bauswein et al. 2018

But: a high frequency on its own may not yet be characteristic for a phase transition

→ unambiguous signature

(→ show that all purely baryonic EoS behave differently)

Signature of 1st order phase transition



from postmerger

from the inspiral

Bauswein et al. 2018

- Tidal deformability measurable from inspiral to within 100-200 (Adv. Ligo design)
- Postmerger frequency measurable to within a few 10 Hz @ a few 10 Mpc (either Adv. Ligo or upgrade)
- Important: “all” purely hadronic EoSs (including hyperonic EoS) follow fpeak-Lambda relation → deviation characteristic for strong 1st order phase transition

Discussion

- Consistency with fpeak-Lambda relation points to

  \- purely baryonic EoS

  \- (or an at most weak phase transition → no strong compactification)

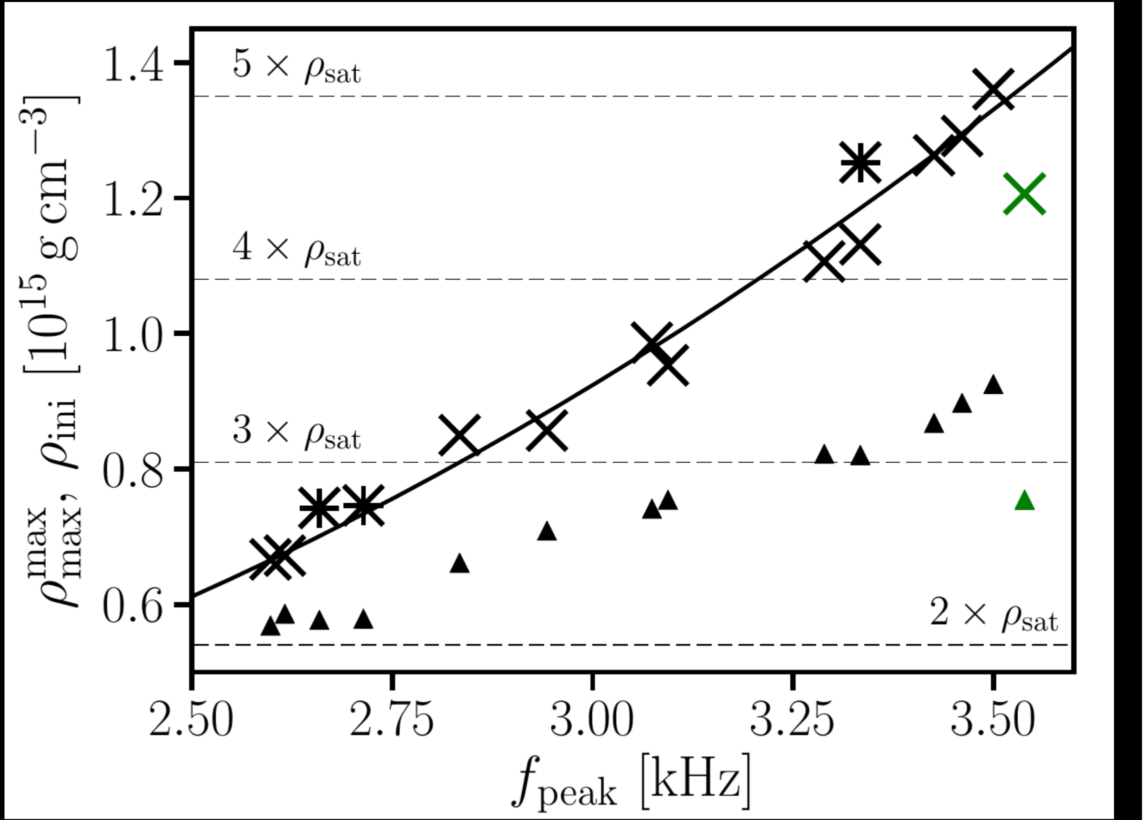
  in the tested (!) density regime

- fpeak also determines maximum density in postmerger remnant

- postmerger GW emission provides complimentary information to inspiral

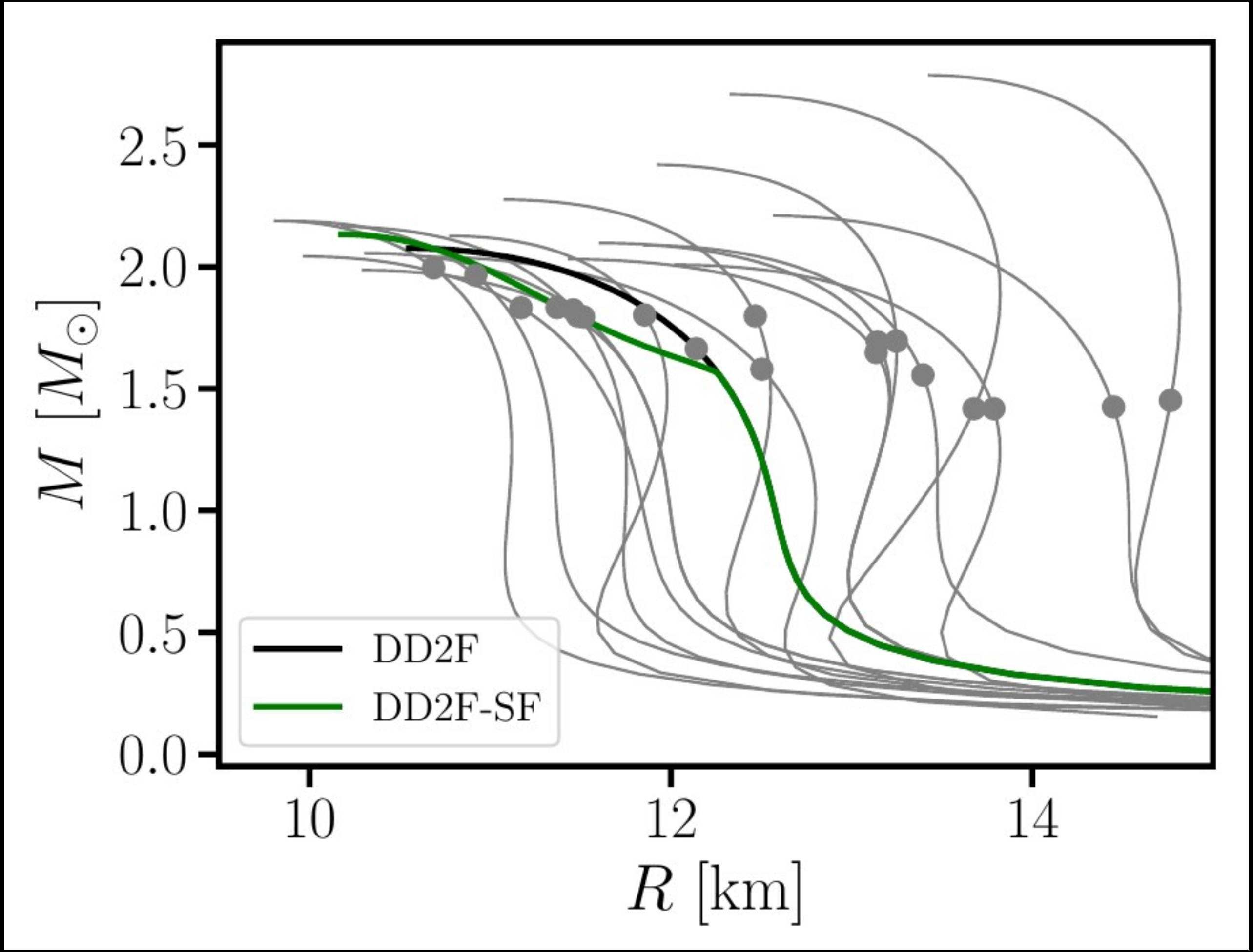
  → probes higher density regime

Bauswein et al. 2018



Probed densities / NS masses

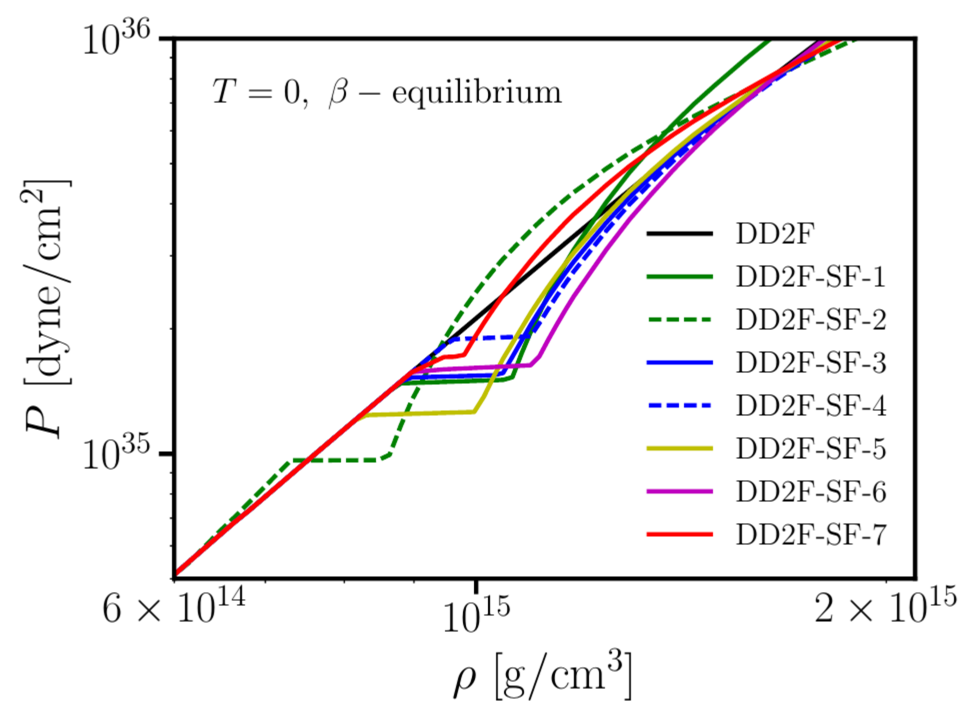
- Dots: NS mass with central density = maximum density during early postmerger evolution

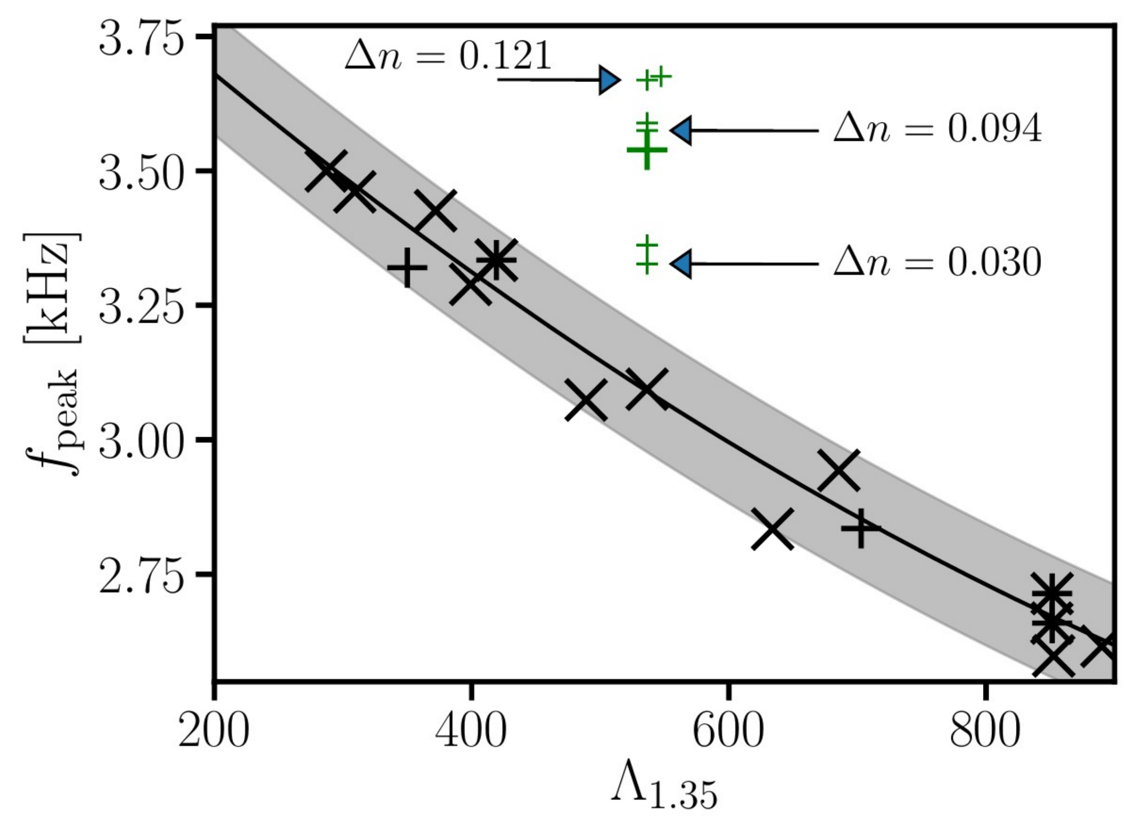


For 1.35-1.35 Msun merger – higher binary masses probe higher densities / NS masses

More models

- Larger density jump → stronger compactification → more significant increase of fpeak (keeping other EoS parameters fixed)





Conclusions

- NS radius must be larger than 10.7 km (very robust and conservative)
- More stringent constraints from future detections
- NS radius measurable from dominant postmerger frequency
- Explicitly shown by GW data analysis
- Threshold binary mass for prompt collapse → maximum mass M_{max}
- Strong 1st order phase transitions leave characteristic imprint on GW (postmerger frequency higher than expected from inspiral)
- Complementarity of inspiral and postmerger phase → postmerger probes higher density regime